

Tunka-Rex: a Radio Antenna Array for the Tunka Experiment

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Abstract. Tunka-Rex, the Tunka radio extension, is an array of 20 antennas at the Tunka experiment close to Lake Baikal in Siberia. It started operation on 08 October 2012. The antennas are connected directly to the data acquisition of the Tunka main detector, a 1 km² large array of 133 non-imaging photomultipliers observing the Cherenkov light of air showers in dark and clear nights. This allows to cross-calibrate the radio signal with the air-Cherenkov signal of the same air showers - in particular with respect to the energy and the atmospheric depth of the shower maximum, X_{max} . Consequently, we can test whether in rural regions with low radio background the practically achievable radio precision comes close to the precision of the established fluorescence and air-Cherenkov techniques. At a mid-term perspective, due to its higher duty-cycle, Tunka-Rex can enhance the effective observing time of Tunka by an order of magnitude, at least in the interesting energy range above 100 PeV. Moreover, Tunka-Rex is very cost-effective, e.g., by using economic Short Aperiodic Loaded Loop Antennas (SALLAs). Thus, the results of Tunka-Rex and the comparison to other sophisticated radio arrays will provide crucial input for future large-scale cosmic-ray observatories, for which measurement precision as well as costs per area have to be optimized. In this paper we shortly describe the Tunka-Rex setup and discuss the technical and scientific goals of Tunka-Rex.

Keywords: radio detection, cosmic rays, air showers, Tunka-Rex

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1. INTRODUCTION

More than 100 years after their discovery [1], cosmic rays are still a useful tool to answer open questions in astro and particle physics. What are the sources of the ultra-high energy cosmic ray particles and how are they accelerated? What is the physics of particle collision in the energy range beyond the center-of-mass energies reachable by LHC? Can the measurement of extensive air-showers give us new insights in the underlying particle physics? Answering these questions requires cosmic-ray observatories for air showers above 10¹⁷ eV, which can surpass the already available measurements in accuracy and statistics. This in turn requires technical developments to improve the detection methods for air showers - aiming at measurement precision and cost-effectiveness at the same time. Digital antenna arrays are thus a promising option to complement air-shower arrays, since radio measurements allow a high-duty cycle and are sup-

posed to provide a high precision at the same time.

The principal feasibility of air-shower radio detection was already shown by historic experiments in the 1950s and 1960s [2, 3]. About 10 years ago, the LOPES [4, 5] and CODALEMA [6] experiments revived the radio technique using digital data acquisition and computer-based analysis procedures. The experience gained in these experiments is now used in state-of-the-art experiments like AERA [7], and Tunka-Rex, the radio extension of the Tunka experiment in Siberia, close to lake Baikal. Tunka-Rex profits also from developments made for AERA, in particular the *Offline* analysis-software framework of the Pierre Auger Observatory [8] will also be used for Tunka-Rex.

The main goal of Tunka-Rex is to test whether radio measurements can indeed provide a measurement precision for the energy and the composition of the primary cosmic rays competitive to the established air-fluorescence and air-Cherenkov techniques. Thus,

Tunka-Rex is built on the same site as the Tunka-133 non-imaging photomultiplier array, which detects the air-Cherenkov light of air-showers, and covers the energy range from about 10^{16} eV to 10^{18} eV [9]. By a cross-calibration of coincident events, the radio precision can be determined with an accuracy limited by the air-Cherenkov precision: about 15% for the energy and 25 g/cm^2 for the atmospheric depth of the shower maximum, X_{max} , which is sensitive to the mass and the interaction of the primary particle.

2. TUNKA-REX SETUP

The Tunka-Rex array consists of 20 antennas with a typical spacing of approximately 200m, and covers an area of about 1 km^2 (figure 1). The pattern follows the hexagonal structure of the Tunka PMT array which consists of 19 clusters with 7 PMTs each, and 6 satellite clusters to increase the area and thus the statistics for high energy events above 10^{17} eV. Each cluster features its own local data-acquisition and trigger. The data from the different clusters are transferred via optical fibers to a central data-acquisition which stores the data. Consequently, merging events from different clusters, data analysis and event reconstruction can be done at any later moment offline.

The radio antennas are connected to the same cluster data-acquisition as the PMTs using free ADC channels on the same electronics board. Thus, the radio signal is read simultaneously with the PMT signal of each shower, which is crucial for performing the cross-calibration between the air-Cherenkov and the radio signal. Moreover, the radio and the air-Cherenkov signal is stored in the same data files which facilitates a hybrid analysis. The technical proof-of-principle was made with a prototype antenna installed in 2009, which detected about 70 radio candidate events in the first year of operation [9, 10].

2.1. Hardware

Since the main goal of Tunka-Rex is to demonstrate a high precision for air-shower observables, the analog radio hardware is optimized for measurement accuracy, accepting a slightly higher detection threshold than otherwise possible. As antenna type we have decided for a SALLA [11] with 120 cm diameter. Although the SALLA provides a signal-to-noise ratio worse compared to other antenna types [12], it has two major advantages: First, it is economic, compact and simple. Second, the gain of the SALLA is relatively independent from environmental conditions. This is important, since the measurement accuracy of pulses with high amplitude is not limited by noise, but by variations of the gain due to en-

TABLE 1. Tunka-Rex technical data.

Area	1 km^2
Number of Antennas	20
Antenna type	SALLA
Frequency range	30 – 80 MHz
Trigger and data acquisition	by Tunka-133 PMT array

vironmental changes, e.g., ground conditions [13, 14].

Like other modern radio arrays, Tunka-Rex features two orthogonal antennas at each position, to study the polarization of the radio signal. Moreover, a reconstruction of the electric field vector and, thus, the total amplitude is possible when the arrival direction of the radio signal is known [15]. Experiments like LOPES, CO-DALEMA, and AERA decided to align their antennas in east-west, respectively north-south direction. This has the advantage of a slightly lower detection threshold, since the radio signal is for most shower geometries mainly east-west polarized [16, 17], but the disadvantage that in many events the north-south component can not be distinguished from noise. For Tunka-Rex we have decided to rotate the antenna alignment by 45° , as it is done at LOFAR [18]. By this we hope to increase the reconstruction accuracy of the electrical field vector, since the signal-to-noise ratio in the two cross-polarized antennas should be roughly equal.

All antennas are connected with RG 213 coaxial cables of about 30m length to the DAQ of the corresponding Tunka cluster. They are placed at a distance of at least 20m to the closest Tunka detector and the DAQ. Thus, possible radio interferences from the PMTs can be distinguished by timing, since the air shower radio pulse is supposed to arrive significantly earlier than any possible disturbance from the PMTs.

In the cluster DAQ, the radio signal is amplified and filtered to a band-width of 30 to 80 MHz. We use steep filters to fulfill the Nyquist sampling theorem, which in our case requires that there any signal or noise at frequencies above 100 MHz is strongly suppressed and thus negligible against the signal in the measurement band-width. Therefore, an accurate reconstruction of the original signal with the bandwidth is possible by up-sampling. We have tested with a climate chamber that the gain of the main amplifier is stable within $\pm 1 \text{ dB}$, and we plan to test the temperature behavior of the whole system including cables and the low noise amplifier of the antenna. For digitizing we use the Tunka electronics described in reference [9]. The signal is digitized with a 12-bit analog-to-digital converter (type AD9430) using a sampling frequency of 200 MHz and a trace length of $5 \mu\text{s}$.

Tunka-Rex will use a reference beacon to monitor and improve the relative timing between different antennas, as LOPES and AERA already do [19, 20]. For this we

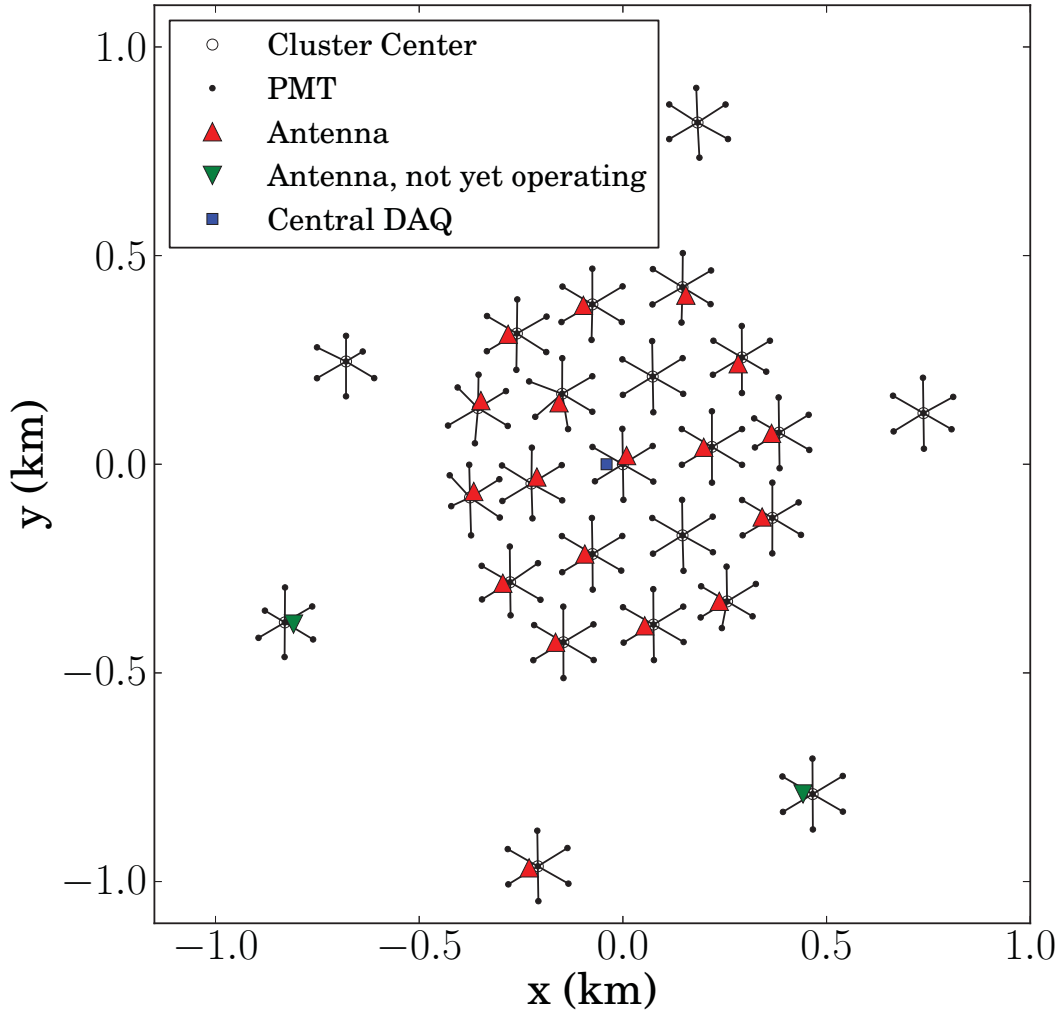


FIGURE 1. Map of the Tunka experiment. The non-imaging photo-multiplier (PMT) array is organized in 25 clusters with 7 PMTs each. Each cluster center features its own digital data acquisition, which is used for the Tunka-Rex antennas, too. The measured data of each cluster are transferred via optical fibers to the central data acquisition (DAQ).

will analyze the phasing of continuous sine waves emitted by the beacon. An accurate relative timing between the different antennas is essential – not only for the reconstruction of the arrival direction. LOPES has shown that the individual antennas can be digitally combined to a radio interferometer, provided that the relative timing is precise and accurate to about 1 ns [19]. Moreover, an accurate timing makes it possible to reconstruct the shape of the radio wavefront which is sensitive to the longitudinal shower development and, thus, to the type and mass of the primary particle [21, 22].

3. GOALS OF TUNKA-REX

The main goal of Tunka-Rex is to test whether the properties of the primary particle can be reconstructed with a precision competitive to other detection technologies. For the arrival direction and the primary energy, LOPES already demonstrated a sufficient precision of 0.65° , respectively 20% [5]. The best way to reconstruct the composition of the primary particles from radio measurements likely is to statistically analyze the distribution of X_{\max} . LOPES demonstrated that X_{\max} can be in principle reconstructed from the radio measurements [23, 24], but to our knowledge no radio experiment has yet demon-



FIGURE 2. Tunka-Rex antenna (SALLA). It is connect to the cluster DAQ (white box) which is close to the one PMT detector.

strated an X_{\max} precision competitive to air-Cherenkov and air-fluorescence measurements. However, the main reason for the bad X_{\max} resolution of LOPES seems to be the high ambient noise level there. In rural areas, the ambient noise level usually is much lower [25], and at least for the rate of background pulses, this is already confirmed by first measurements at Tunka. Thus, for Tunka-Rex we expect a better X_{\max} precision, which we will check by a cross-calibration with the radio and air-Cherenkov hybrid events.

Another goal of Tunka-Rex is to improve the understanding of the radio emission by air showers. We will study the properties of the radio signal, in particular the lateral distribution, the shape of the wavefront, the polarization of the radio signal, and the frequency spectrum. Thus, we can cross-check results from other experiments. E.g., historic experiments pointed towards an exponential lateral distribution function (LDF) [3], which is also used by LOPES [26], but newer results obtained at AERA indicate that a Gaussian LDF is more appropriate to fit a large distance range [15]. The wavefront shape so far has only been studied by LOPES and was found to

be approximately conical [22]. The polarization seems to be predominately determined by the geomagnetic effect [27, 16], but is influenced by the Askaryan effect, too [28, 29]. The shape of the frequency spectrum is not yet well known. It is falling towards higher frequencies [30], and its slope seems to be sensitive to the longitudinal shower development [31], which provides another method for X_{\max} reconstruction.

Moreover, we can use Tunka-Rex to test models for the radio emission. A recent comparison of LOPES measurements with REAS 3.11 simulations [32, 33] indicates that REAS 3.11 can describe measured radio events – in contrast to some other radio simulation codes not including the refractive index of air, which changes the coherence conditions for the radio emission [34]. Due to the lower ambient noise level, we expect that Tunka-Rex can test models to a higher precision than LOPES. In addition, we plan to compare Tunka-Rex measurements to the predictions of several radio simulation codes, e.g., the one described in reference [35], which will help to further improve the understanding of the physics processed behind the radio signal from air showers.

4. CONCLUSION

Tunka-Rex is based on the success of many other radio arrays for air shower measurements. It makes use of the world-unique opportunity to perform a cross-calibration between air-Cherenkov and radio measurements of the same cosmic ray events. Thus, it will test the expectation whether radio measurements in rural areas really allow a reconstruction precision for the energy and X_{\max} similar to the precision of the air-Cherenkov measurements. If so, Tunka-Rex can enhance the duty cycle of Tunka for high energy events by an order of magnitude, since it can measure also during days and during almost any weather conditions (except close-by thunderstorms). For this we foresee to trigger Tunka-Rex by the planned scintillator extension of Tunka. Finally, comparing the results of Tunka-Rex with other next generation experiments like AERA will reveal whether a sufficient understanding of the radio signal is already achieved, and under which conditions it makes sense to extend future cosmic ray observatories by digital antenna arrays.

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